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The VANDELS survey: the role of ISM and galaxy physical properties on the escape of Ly α emission in $z \sim 3.5$ star-forming galaxies [★]

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ABSTRACT

Aims. We wish to investigate the physical properties of a sample of Ly α emitting galaxies in the VANDELS survey, with particular focus on the role of kinematics and neutral hydrogen column density in the escape and spatial distribution of Ly α photons.

Methods. From all the Ly α emitting galaxies in the VANDELS Data Release 2 at $3.5 \lesssim z \lesssim 4.5$, we select a sample of 52 galaxies which also have a precise systemic redshift determination from at least one nebular emission line (HeII or CIII). For these galaxies, we derive different physical properties (stellar mass, age, dust extinction and star formation rate) from spectral energy distribution (SED) fitting of the exquisite multi-wavelength photometry available in the VANDELS fields, using a state-of-the-art spectral modeling tool, BEAGLE and the UV β slope from the observed photometry. We characterize the Ly α emission in terms of kinematics, EW, FWHM and spatial extension and then estimate the velocity of the neutral outflowing gas. Thanks to the ultra-deep VANDELS spectra (up to 80 hours on-source integration) this can be achieved for individual galaxies, without relying on stacks. We then investigate the correlations between the Ly α properties and the other measured properties, to study how they affect the shape and intensity of Ly α emission.

Results. We reproduce some of the well known correlations between Ly α EW and stellar mass, dust extinction and UV β slope, in the sense that the emission line appears brighter in lower mass, less dusty and bluer galaxies. We do not find any correlation with the SED-derived star formation rate, while we find that galaxies with brighter Ly α tend to be more compact both in UV and in Ly α . Our data reveal a new interesting correlation between the Ly α velocity and the offset of the inter-stellar absorption lines with respect to the systemic redshift, in the sense that galaxies with larger inter-stellar medium (ISM) out-flow velocities show smaller Ly α velocity shifts. We interpret this relation in the context of the shell-model scenario, where the velocity of the ISM and the HI column density contribute together in determining the Ly α kinematics. In support to our interpretation, we observe that galaxies with high HI column densities have much more extended Ly α spatial profiles, a sign of increased scattering. However, we do not find any evidence that the HI column density is related to any other physical properties of the galaxies, although this might be due in part to the limited range of parameters that our sample spans.

Key words. Galaxies: Star-Forming Galaxies, Ly α emission

1. Introduction

The Ly α emission line plays a fundamental role in the investigation of several astrophysical problems, from detecting very

[★] Based on data obtained with the European Southern Observatory Very Large Telescope, Paranal, Chile, under Large Program 194.A-2003(EK).

distant sources, to inferring galaxy physical parameters. Understanding the mechanisms that regulate this strong emission, is critical for interpreting the galaxy populations during the reionization epoch, since the most luminous Ly α emitters (LAE) are frequent sources among the population of high- z galaxies (e.g. Stark et al. 2010; Pentericci et al. 2009; Stark et al. 2017). In addition, it is believed that the mechanisms that regulate the escape of Ly α photons are also responsible for the escape of LyC photons (Verhamme et al. 2017; Steidel et al. 2018; Marchi et al. 2018). Constraining these processes is therefore fundamental to investigate the ionizing radiation from high redshift star-forming galaxies.

Since Ly α photons are mainly produced in HII regions by recombination processes, the line strength is related to the UV radiation emitted by young stars in galaxies, and so to the on-going star formation. The escape of Ly α photons, depends, however, on a variety of physical processes. The interpretation from theoretical studies is that these photons are scattered by the neutral gas and absorbed by dust in the inter-stellar medium (ISM) (e.g. Verhamme et al. 2006; Gronke & Dijkstra 2016) and that these processes have a fundamental role in shaping the Ly α line profile, which is usually characterized by a single peak, redshifted with respect to the systemic redshift, plus a blue-ward emission. In many LAEs, however, we can clearly distinguish two different peaks on the Ly α spectrum (e.g. Yamada et al. 2012; Verhamme et al. 2017; Matthee et al. 2018), red-ward and blue-ward of the systemic position, with the red peak more commonly dominant with respect to the blue one. According to models, the shift of the red peak, the intensity of the blue peak and its distance from the red peak critically depend on the neutral hydrogen column density ($N(\text{HI})$) and the dust content of galaxies (e.g. Verhamme et al. 2015).

The physical properties of Ly α emitting galaxies and the relation between the Ly α emission and galaxy physical properties have been extensively investigated by several groups during the last decade, using observations of large samples of galaxies at different redshifts. In particular, a bright Ly α emission appears to be associated with galaxies showing lower UV luminosities and lower metallicities (e.g. Shapley et al. 2003; Reddy et al. 2006) and lower dust content. It is also believed that lower stellar masses and lower SFRs characterize the brightest LAEs, even if for these two quantities a general consensus has not been yet achieved (see e.g. Pentericci et al. 2010; Kornei et al. 2010; Hathi et al. 2016; Du et al. 2018). While some authors advocate a scenario in which Ly α emitting galaxies represent early stages in galaxy formation (e.g. Finkelstein et al. 2007; Cowie et al. 2011), others find that these galaxies actually span a wide range of physical properties (e.g. Pentericci et al. 2009; Finkelstein et al. 2009; Kornei et al. 2010; Finkelstein et al. 2015).

The recently completed spectroscopic survey VANDELS (a VIMOS survey of the CANDELS UDS and CDFS fields, Pentericci et al. 2018; McLure et al. 2018) has obtained spectra of star forming galaxies spanning a wide range of redshifts and with unprecedented depth. Thanks to this new data-set we can analyze the Ly α line properties in extremely deep and moderate resolution spectra for a large sample of $z \sim 3.5$ galaxies, and infer the galaxy properties that most correlate with the intensity and kinematics of this line. In particular, we aim to assess the role of ISM kinematics and neutral hydrogen column density in the escape and distribution of Ly α photons. The paper is organized as follows. In Section 2 we describe the criteria that we used to select the sample. In Section 3 we describe the methods used to measure the spectral properties from the deep VAN-

DELS spectra and the physical properties from SED fitting of the deep multi-band imaging available. In Section 4 we describe how our sample relates with a more generic sample of galaxies with Ly α emission also from the VANDELS survey. In Section 5 we present the results of this analysis and show the correlations that we found between Ly α and galaxies properties. In Section 6 we finally discuss the main result, that is the correlation between the Ly α peak shift and the velocity of the gas inside the ISM, and interpret this relation in the context of the shell model, assessing the role of the gas kinematics and HI column density in the Ly α photons escape. Throughout the paper we adopt the Λ cold dark matter (Λ -CDM) cosmological model ($H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$ and $\Omega_\Lambda = 0.7$) and express all magnitudes in the AB system. Since the VANDELS spectra are calibrated in air and not in vacuum, we use air wavelengths to fit the position of the different lines. We also use positive values for the EWs of emission lines and negative for absorption lines. Our EWs are always rest-frame values.

2. Sample selection

To study the Ly α emission properties of high redshift star-forming galaxies, we exploited the extremely rich and deep VANDELS data-set. In this section we first describe the VANDELS survey and then the criteria that we applied to select our sample.

2.1. The VANDELS survey

VANDELS, a VIMOS survey of the CANDELS UDS and CDFS fields (PI L. Pentericci and R. McLure) is an ESO public spectroscopic survey mainly targeting high-redshift galaxies at $z > 2.5$ and specifically designed to be the deepest-ever spectroscopic survey of the high-redshift Universe. The survey is presented in two companion papers, McLure et al. (2018) that provides an overview of the survey strategy and target selection, and Pentericci et al. (2018) that is focused on the observations and the first data release. VANDELS exploited the VIMOS multi-object spectrograph to obtain ultra-deep optical spectra of ~ 2100 high-redshift galaxies within the redshift interval $1 \lesssim z \lesssim 7$, covering a total area of 0.2 deg^2 centered on the CANDELS-UDS (Ultra Deep Survey) and CANDELS-CDFS (*Chandra* Deep Field South) fields (Grogin et al. 2011; Koekemoer et al. 2011). Thanks to the ultra-long exposure time strategy, each target has been observed from a minimum of 20 hours to a maximum of 80 hours, using the red medium-resolution grism of VIMOS, that covers a wavelength range of $4800 \text{ \AA} \leq \lambda \leq 10000 \text{ \AA}$ with an average resolution of 580. To select VANDELS sources, in the CANDELS/HST areas (CDFS-HST and UDS-HST) the photometric catalogs based on H_{160} -band detections and provided by the CANDELS team (Galametz et al. 2013; Guo et al. 2013) were adopted. Specifically, the CDFS-HST catalog is based on 17 photometric broadband filters (Guo et al. 2013), while the UDS-HST catalog includes photometry in 19 broadband filters (Galametz et al. 2013). However, since the VIMOS footprint is wider than the one of HST, many VANDELS sources fall outside the CANDELS footprints but are still covered by ground-based and Spitzer IRAC imaging. The two photometric catalogs used for the sources in the wide-field regions were generated by McLure et al. (2018) using the publicly available imaging in 12 filters for UDS-GROUND and 17 filters for CDFS-GROUND, spanning the range from U to K band in both cases. These photometric catalogs, not only will be combined with the

spectroscopic data to create these unique data-set to study the high redshift galaxies properties and evolution, but have also allowed the VANDELS team to carefully pre-select the VANDELS targets on the basis of their accurate photometric redshift measures (McLure et al. 2018). $\sim 85\%$ of the galaxies were selected to be at $z \geq 3$, the majority of which being Lyman Break Galaxies (LBG), which are star-forming galaxies selected on the basis of their colors red-ward and blue-ward of the observed-frame Lyman Limit. The mask were prepared using the VIMOS mask preparation software (Bottini et al. 2005) that is distributed by ESO. The VANDELS spectroscopic data were reduced using the fully automated pipeline Easy-Life (Garilli et al. 2012), an updated version of the algorithms and dataflow of the original VIPGI system (Scodeggio et al. 2005), that generated wavelength- and flux-calibrated spectra directly from the raw data. Finally, the spectroscopic redshifts have been estimated using the Pandora software package, within the EZ environment (Garilli et al. 2010). The same procedure adopted in VUDS was applied (see Le Fèvre et al. 2015), i.e. each spectrum was analysed by two different team members. In addition a final independent check by the two PIs was carried out. At present, the same VUDS quality flags (see Le Fèvre et al. 2015) are used to assess the spectroscopic redshift reliability.

The second VANDELS data release, became public at the beginning of October 2018 and contains the spectra and redshifts of 1339 objects (557 in CDFS and 762 in UDS) with more than 200 spectra with ultra-deep 80 hours exposures¹.

2.2. The sample

We selected all the objects in the VANDELS data release 2 which showed Ly α in emission and had a secure spectroscopic redshift with VANDELS reliability flags 3 or 4 (corresponding to a probability greater than 95% for the spectroscopic redshift to be correct, see Pentericci et al. 2018). This sample, which we will call the *parent sample*, contains 305 star-forming galaxies in the redshift range $3 \lesssim z \lesssim 4.5$. 146 of these galaxies are in the CDFS field and 159 are in UDS. 154 (78 in the CDFS and 76 in UDS) fall within the CANDELS regions (Grogin et al. 2011; Koekemoer et al. 2011), and therefore are covered by very deep multi-wavelength HST imaging, while the remaining 151 (68 and 83 in the CDFS and UDS, respectively) are in the wide-field areas of the VANDELS fields and mainly benefit from ground-based imaging.

Since we want to study the kinematics of the Ly α emission, which gives us important information about the ISM and the HI column density as shown by Verhamme et al. (2015), we need a precise measure of the velocity of the Ly α line with respect to the systemic redshift of the galaxy. Therefore, we restricted our sample to those galaxies for which we could estimate the systemic redshift in a reliable way. The systemic redshift (z_{sys}) is ideally defined by the positions of photospheric absorption lines, that are generated by the absorption of the stellar radiation by the most external layers of the star. The most important photospheric absorption lines in the UV regime for highly star forming galaxies would be the OIV $\lambda 1343\text{\AA}$, SiII $\lambda 1417\text{\AA}$ and SV $\lambda 1500\text{\AA}$ (Talía et al. 2012). However, these lines are either too weak or even not detected in our ultra-deep spectra. As an alternative we can use non-resonant nebular emission lines which originate from the photoionization of nebular regions by the radiation from very massive O and B stars and can be therefore used to trace the

systemic redshift. The only nebular emission line that enters the wavelength range covered by the VANDELS spectra at redshift ~ 4 is the CIII] $\lambda 1909\text{\AA}$ emission. This is a semi-forbidden line at $\lambda = 1907.7\text{\AA}$ and it is known to be a doublet at a rest-frame vacuum wavelength of 1906.68\AA and 1908.68\AA (note that, since the VANDELS spectra are air-calibrated, we use in this paper the corresponding air wavelengths 1906.05\AA and 1908.05\AA). There is however another emission line that often appears in our spectra that can trace the systemic redshift, namely the HeII at rest-frame vacuum wavelength 1640\AA (1639.83\AA in air). This feature can be produced either by Wolf-Rayet stars (WR), that are thought to represent a stage in the evolution of O-type stars with masses $> 20M_{\odot}$, or by young star clusters with intense star formation (e.g. Talía et al. 2012; Cassata et al. 2013). Recently, Schaerer et al. (2019) also proposed that high-mass X-ray binaries could be the main source of nebular HeII emission in low-metallicity star-forming galaxies.

The line can show a broad profile because of the strong winds powered by the intense star formation, or can appear as a nebular narrow-line as the strong UV emission from the stars photoionizes the surrounding medium. Because of the possible effects of the winds, this line is not always a good tracer of the systemic redshift. However the results presented in this paper do not change if we exclude the sources whose systemic redshift was derived using HeII $\lambda 1640\text{\AA}$, suggesting that at least for our galaxies, this line provides a good estimate of the systemic redshift.

We therefore selected from the parent sample all those galaxies that showed a CIII] $\lambda 1909\text{\AA}$ and/or HeII $\lambda 1640\text{\AA}$ emission lines, to derive the systemic redshift. We only included objects where the CIII] $\lambda 1909\text{\AA}$ (or HeII $\lambda 1640\text{\AA}$) were reasonably clear of skyline residuals and had a $S/N \geq 3$, to allow a precise measurement of the line peak (see below). With these criteria we selected 60 star-forming galaxies in a redshift range of $3 \lesssim z \lesssim 4.5$.

From this sample we further excluded few objects whose Ly α was too noisy to give a precise estimate of the peak shift and the EW, and those where residuals from skylines fell close to the Ly α emission line, preventing us from carrying out the measures in a reliable way. We also excluded two objects with intense CIV $\lambda 1549\text{\AA}$ emission, that could indicate the presence of an AGN.

The final *selected sample* contains 52 sources. 29 (25) of these are in the CDFS (UDS) field. 30 (18 in the CDFS and 12 in UDS) are in the CANDELS regions (Grogin et al. 2011; Koekemoer et al. 2011), and therefore are covered by very deep multi-wavelength HST imaging, while the remaining 24 (11 and 13 in the CDFS and UDS respectively) are in the wide-field areas of the VANDELS fields. Through out the paper, we refer to this group of galaxies as simply the *selected sample*.

The distributions in redshift and z-band magnitude (that probes the rest-frame UV continuum at $\sim 1800\text{\AA}$ for all sources) of the *parent sample* and the *selected sample* are shown in Fig. 1 in yellow and green, respectively. The median redshift and the median z-band magnitude are $z_{\text{med}} = 3.47$ and $z_{\text{mag,med}} = 25.4$, for the *selected sample* and $z_{\text{med}} = 3.47$ and $z_{\text{mag,med}} = 25.5$, for the *parent sample*. The two samples have very similar UV continuum luminosity distributions. However, since the selected sample mostly contains galaxies that also show CIII] $\lambda 1909\text{\AA}$ emission, we expect this sample to be in part biased to brighter Ly α luminosities, given that these two emissions appear to be correlated from several studies (e.g. Shapley et al. 2003; Stark et al. 2014; Maseda et al. 2017).

¹ The VANDELS second data release can be found at <https://www.eso.org/sci/publications/announcements/sciann17139.html>

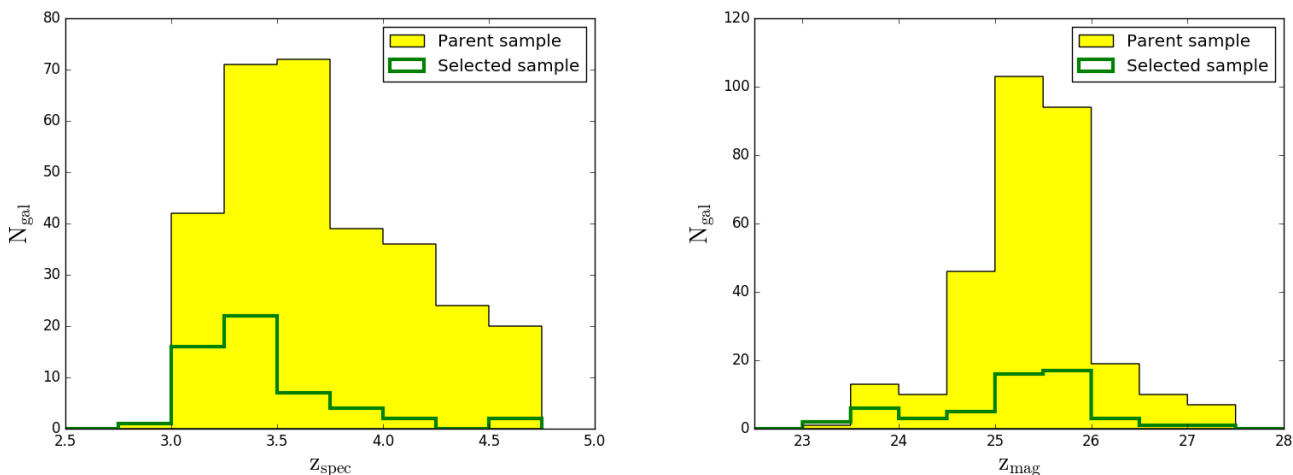


Fig. 1. Redshift distribution (*Left panel*) and z magnitude distributions (*Right panel*) of the sources in the *selected sample* (green empty histogram) and in the *parent sample* (yellow histogram).

Finally, we remark that although our samples have been selected on the presence of Ly α emission, only a fraction could be strictly considered LAEs if we apply the traditional threshold of $EW(Ly\alpha) > 25\text{\AA}$. In our parent sample, 26% of all galaxies has $EW(Ly\alpha) \geq 25\text{\AA}$ in agreement with the statistics at this redshift (Shapley et al. 2003; Stark et al. 2010).

3. Method

For each galaxy in our samples, the VANDELS dataset provides not only a deep spectrum but also deep optical, near-IR and Spitzer photometry. To best exploit this unique legacy dataset, we therefore used the optical spectra to characterize the Ly α emission, both spectrally and spatially, the inter-stellar absorption lines and the CIII] $\lambda 1909\text{\AA}$ properties, while we used the VANDELS photometric catalogs (McLure et al. 2018) to derive the physical parameters. We note that for the parent sample we only measured the Ly α EW from the 1-dimensional spectra and the physical properties from the available photometry.

3.1. Spectroscopic measures

From the VANDELS spectra we measured the following properties.

– The systemic redshift (z_{sys})

We evaluated the systemic redshift from the CIII] $\lambda 1909\text{\AA}$ line or from the HeII $\lambda 1640\text{\AA}$ line. In the majority of galaxies with CIII] $\lambda 1909\text{\AA}$ emission, the two carbon lines appear blended; for these and for the objects with HeII $\lambda 1640\text{\AA}$ emission, we derived the systemic redshift from the observed-frame spectra by fitting a single Gaussian profile assuming that the continuum has a well-defined slope (linear fit). In the minimization procedure we took into account the VANDELS noise spectrum, that contains both the observational and instrumental uncertainties. While fitting the continuum, we excluded the wavelength regions with absorption lines or where strong residuals from sky lines were present, because these regions could alter the slope of the fit. This procedure was done by visual inspection in each spectrum. We finally estimated the systemic redshift from the best-fit mean of the Gaussian profile. We show an example of a fit of

the CIII] $\lambda 1909\text{\AA}$ in the bottom right panel of Figure 2. The blue line is the VANDELS spectrum, the magenta range is the part of the spectrum that we are fitting and the dark blue line is the best fit. For the few objects that clearly show the two peaks of the CIII] $\lambda 1909\text{\AA}$ line due to the better S/N, we used the same procedure but including two Gaussian profiles, and we evaluated the systemic redshift directly from the fit imposing the position of the two mean peaks to be at $1906.68(1+z)\text{\AA}$ and $1908.68(1+z)\text{\AA}$. We did not apply any restriction to the CIII] doublet ratio during the fitting procedure, but the ratios that we obtain have values in the range $\sim 1 - 1.5$ that are typical values for the electronic densities of local star-forming regions (see e.g. Osterbrock & Ferland 2006; Quider et al. 2009). To derive the errors on these measures we used the equation given by Lenz & Ayres (1992) for a Gaussian fit.

– The Ly α redshift and Ly α velocity shift ($Ly\alpha_{\text{shift}}$)

The Ly α redshift was evaluated as the peak of the Ly α line i.e. at the position where the flux has a maximum. Contrary to the CIII] emission we did not attempt a Gaussian fit of the Ly α line, since its shape is known to be asymmetric at high redshift (e.g. Shimasaku et al. 2006). For the few spectra (5 out of 52) where we could clearly identify a double peak profile of the Ly α , we took the peak position of the red peak. We then evaluated the Ly α velocity shift as the velocity difference between the Ly α redshift and the systemic redshift. The error has been evaluated from the usual propagation formula using the error on z_{sys} and the error for the Ly α wavelength position evaluated in a similar way.

In some spectra the wavelength range around the Ly α was affected by noise, e.g., due to residuals from skylines, and therefore it was not possible to derive an accurate peak position of the Ly α line. We decided to keep these spectra to derive the other spectral parameters, such as $EW(Ly\alpha)$ (see below). In the end we could measure the Ly α velocity shift only for 46 galaxies out of 52. We show for reference the Ly α line profile of the galaxy CDFS122764 in the bottom left panel of Figure 2. Here, the Ly α velocity shift is the velocity difference between the peak of the line (dotted line) and the green line that represents the position of the line at systemic redshift.

– The Ly α FWHM

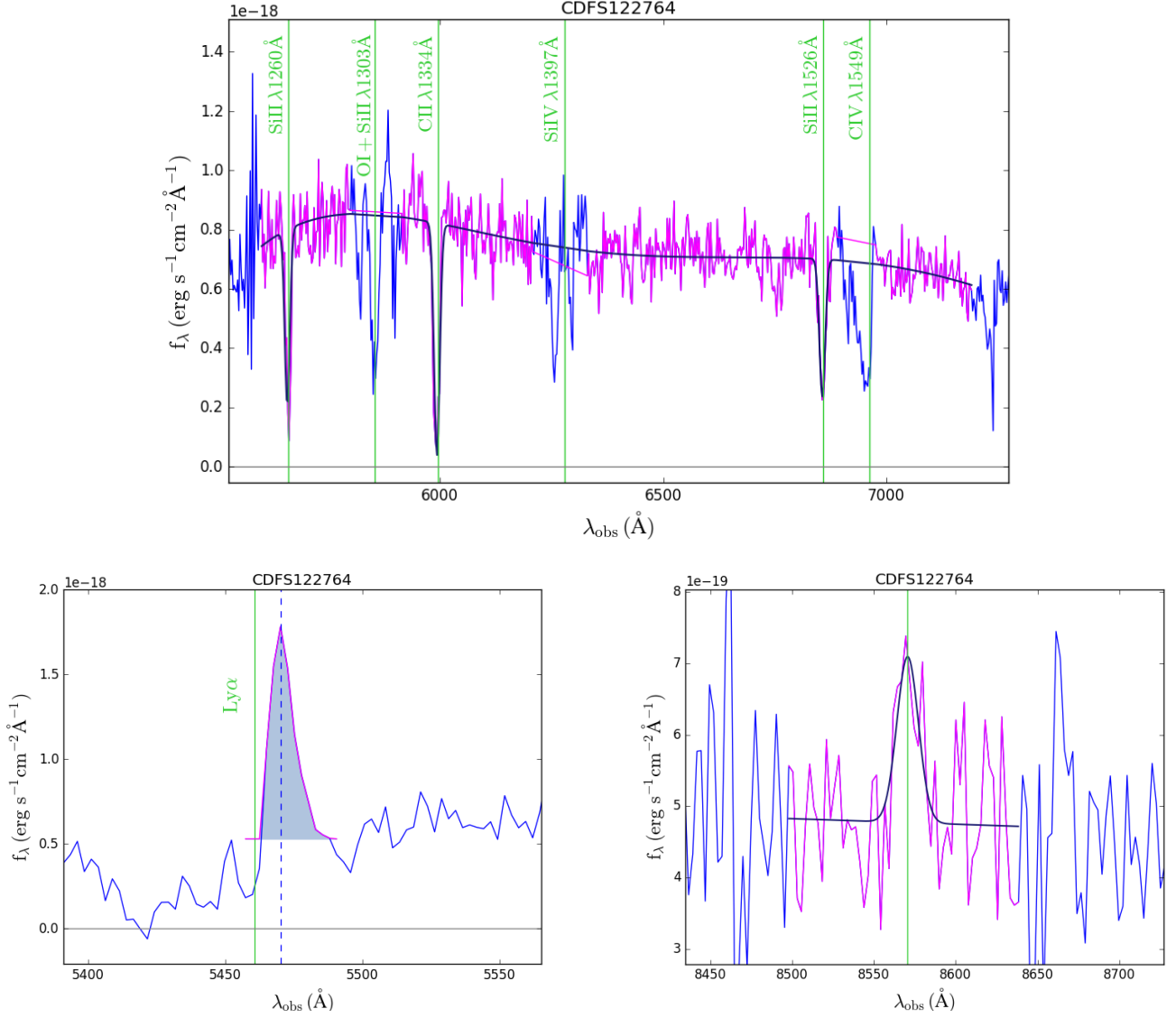


Fig. 2. The different panels show different regions of the spectrum of the object CDFS122764 at $z = 3.49$ that we show for reference. The *upper panel* shows the UV continuum between $\sim 1250 - 1750 \text{ Å}$ rest-frame, where the most important absorption features are clearly visible, the *lower left panel* shows the Ly α line profile and the *right lower panel* shows the CIII] $\lambda 1909 \text{ Å}$ emission. The magenta parts of the spectrum are the regions that we used to perform the fits. The dark blue lines in the *upper panel* and the *lower right panel* are the best fits and the shaded region in the *left lower panel* represent the normalized integrated flux of the Ly α that we used to evaluate the EW of the line. The vertical green lines show the positions of the lines at rest-frame air wavelength. The blue vertical dashed line in the *left lower panel* indicates the position of the Ly α peak.

For the same 46 objects for which we measured the Ly α peak position, we also evaluated the FWHM of the line, directly taking the wavelength interval at half of the flux of the peak of the line. Again we prefer a direct measure of the FWHM rather than a Gaussian fit since in many cases the line is clearly asymmetric and the Gaussian is not a good representation of its shape. We estimated the error using the formula in Lenz & Ayres (1992).

– The Ly α and CIII] $\lambda 1909 \text{ Å}$ EWs

To estimate the Ly α equivalent width we first normalised the line flux. We used as continuum the mean flux in a spectral range between the Ly α and the SiII at 1260 Å that is the first strong absorption line red-ward the Ly α that we can

clearly see in the VANDELS spectra². We then evaluated the EW(Ly α) from the integrated normalised flux. To estimate the error on the EW(Ly α) we used equation 13 in Ebbets (1995). We show for reference the Ly α line profile of the galaxy CDFS122764 in the bottom left panel of Figure 2, where the shaded region represents the integrated flux of the line.

The EW(CIII]) has been evaluated directly from the fit used to derived the systemic redshift, normalizing the line flux with the best fit line of the continuum. As for the EW(Ly α) we used the Ebbets (1995) formula to derive the errors, using as continuum uncertainty the standard deviation in a wave-

² We did not include possible effects from HI absorption around the Ly α in our estimate of continuum. This might result in a slight underestimate of the EW.

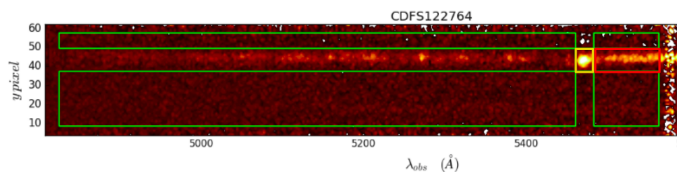


Fig. 3. The two dimensional spectrum of the object CDFS122764 that we show to visualize the method that we used to evaluate the $\text{Ly}\alpha$ and the UV extents. The green big rectangles show the spatial regions of the two dimensional spectrum that we used to estimate the background, while the yellow (red) square around the $\text{Ly}\alpha$ (UV continuum) is the region where we collapsed the spectrum.

length range of 100\AA observed-frame red-ward the CIII $\lambda 1909\text{\AA}$ line.

– *The LIS redshift (z_{LIS}) and velocity shift (IS_{shift})*

The low-ionization absorption lines (LIS) are among the strongest lines detectable in UV rest-frame spectra of star-forming galaxies. In particular the strongest lines that we can observe in the majority of our spectra are the SiII $\lambda 1260\text{\AA}$, CII $\lambda 1334\text{\AA}$, SiII $\lambda 1526\text{\AA}$, FeII $\lambda 1608\text{\AA}$ and the Al III $\lambda 1679\text{\AA}$. In high redshift star forming galaxies, whose UV continuum is dominated by young stars, these lines are mainly produced by the absorption of the stellar radiation by the ISM and are a fundamental tool to study its properties.

We used the velocity shift of the LIS to probe the velocity of the ISM. We chose to use only the lines with the highest S/N in our spectra: SiII $\lambda 1260\text{\AA}$, CII $\lambda 1334\text{\AA}$ and SiII $\lambda 1526\text{\AA}$. We evaluated the redshift of these lines using a combined fit of all the three features, even though in some cases we excluded the lines that did not have a sufficient S/N (less than 3). We fitted the continuum with a fifth order polynomial function, plus three Gaussian profiles to fit the lines (or less if we could not fit all the lines) to derive z_{LIS} , imposing the line centers at $1259.33(1+z)$, $1333.95(1+z)$ and $1526.13(1+z)$, respectively. We again excluded the wavelength regions where other absorption lines were visible or where strong residuals from sky lines were present.

We then evaluated the LIS velocity shifts as the velocity difference between the LIS redshift and the systemic redshift. The error has been evaluated from the error propagation formula using the error on z_{sys} and the error for the LIS redshift. We could estimate the IS_{shift} only for 27 galaxies because the S/N of the LIS for the remaining sources was not sufficient to evaluate the IS velocity. We show the spectral continuum of the galaxy CDFS122764 in the top panel of Figure 2 for reference. The blue line is the VANDELS spectrum and the magenta line is the part of the spectrum that we used for fitting the LIS and the continuum. For this object we could fit all the three LIS, and we excluded from the fit the strong absorption lines (for instance OIdb $\lambda 1303\text{\AA}$, the SiIV+OIV $\lambda 1397\text{\AA}$ and the CIV $\lambda 1549\text{\AA}$) that would have altered the shape of the continuum.

– *The $\text{Ly}\alpha$ and UV spatial extensions*

The $\text{Ly}\alpha$ spatial extent was evaluated directly from the two-dimensional spectrum of the sources. To estimate the background, we selected four spatial windows on the 2d spectrum, as shown in Figure 3, which do not contain the continuum nor the $\text{Ly}\alpha$ residuals that can affect the regions below and above the $\text{Ly}\alpha$ emission in the 2d spectrum. We then separately collapsed the two bottom and upper windows and the region that contains the $\text{Ly}\alpha$ and evaluated the mean value for each y-pixel. We then fitted it with a Gaussian profile us-

ing a straight line to fit the background and finally evaluated the $\text{Ly}\alpha_{\text{ext}}$ as the FWHM of the best fit. To determine the uncertainties in these measurements, we evaluated not only the error on the Gaussian fit as in Lenz & Ayres (1992), but also included the error on the background which we evaluated performing Monte Carlo simulations by varying the value of each y-pixel within the standard deviation of the background and fitting each time with a Gaussian profile. The final uncertainties are computed by combining the two errors in quadrature, assuming that they are independent. We note that the two different errors are comparable on average. We evaluated the UV spatial extension in the same way we measured the $\text{Ly}\alpha$ extent but using for the continuum the spectral region between the two spatial windows red-ward the $\text{Ly}\alpha$ as shown with a red rectangle in Figure 3³.

Our measurements of the $\text{Ly}\alpha$ and UV sizes cannot be taken as absolute estimates of these quantities since we did not deconvolve them by the PSF of the observations. However, this is not a crucial issue for the present analysis since we are only interested in relative values and given that the observations were taken under very similar average conditions over many separate nights, we can assume that the PSF affects all the sources in the same way (see Pentericci et al. 2018).

3.2. Physical parameters

3.2.1. M^* , SFR, E(B-V) and mass-weighted galaxy age

We performed spectral energy distribution (SED) fitting to derive the physical parameters of the sources in our sample. For the objects that fall within the CANDELS CDFS and UDS footprints, we used the CANDELS official photometric catalogs (Galametz et al. 2013; Guo et al. 2013), while for the galaxies in the wide-field areas, we used the photometric catalogs described in McLure et al. (2018). Both catalogs span a wavelength range that goes from the U band to the K band observed-frame.

To perform the SED fitting we used BEAGLE (for BayESian Analysis of GaLaxy sEds, Chevallard & Charlot 2016), a new-generation tool that allows one to model spectroscopic and photometric galaxy observations with a self-consistent physical model. We used BEAGLE with the most recent prescriptions of Charlot & Bruzual (in prep) describing the emission from stars and the models from Gutkin et al. (2016) to account for emission from photoionized gas. The redshift was fixed at the systemic value that we evaluated from the VANDELS spectra for the objects in the selected sample, while we used the official VANDELS spectroscopic redshift for the galaxies in the parent sample. We used the Chabrier (2003) initial mass function and a Calzetti et al. (2000) attenuation law, since it was recently shown by Cullen et al. (2018) that VANDELS galaxies at $3 < z < 4$ are consistent with this law. We used a constant star formation history (SFH) and fitted the metallicity using a Gaussian prior centered at $0.14 Z_{\odot}$ ⁴ with a σ of $0.07 Z_{\odot}$. The value of $0.14 Z_{\odot}$ corresponds to the average metallicity of our sample evaluated applying the method described in Sommariva et al. (2012) on the stacked spectrum of all the galaxies in the selected sample. Briefly, this method exploits spectral indices that are present in the UV spectra of star-forming galaxies in the range between $\sim 1300\text{\AA}$ and 1600\AA rest-frame, whose EWs depend only on

³ The typical spatial windows are $\sim 25\text{\AA}$ observed-frame \times 30 pixels for the $\text{Ly}\alpha$ region and $\sim 150\text{\AA}$ observed-frame \times 10 pixels for the UV continuum. We further checked these regions objects by objects visually inspecting the 2-dimensional spectra.

⁴ BEAGLE adopts a solar metallicity of $Z_{\odot} = 0.01524$

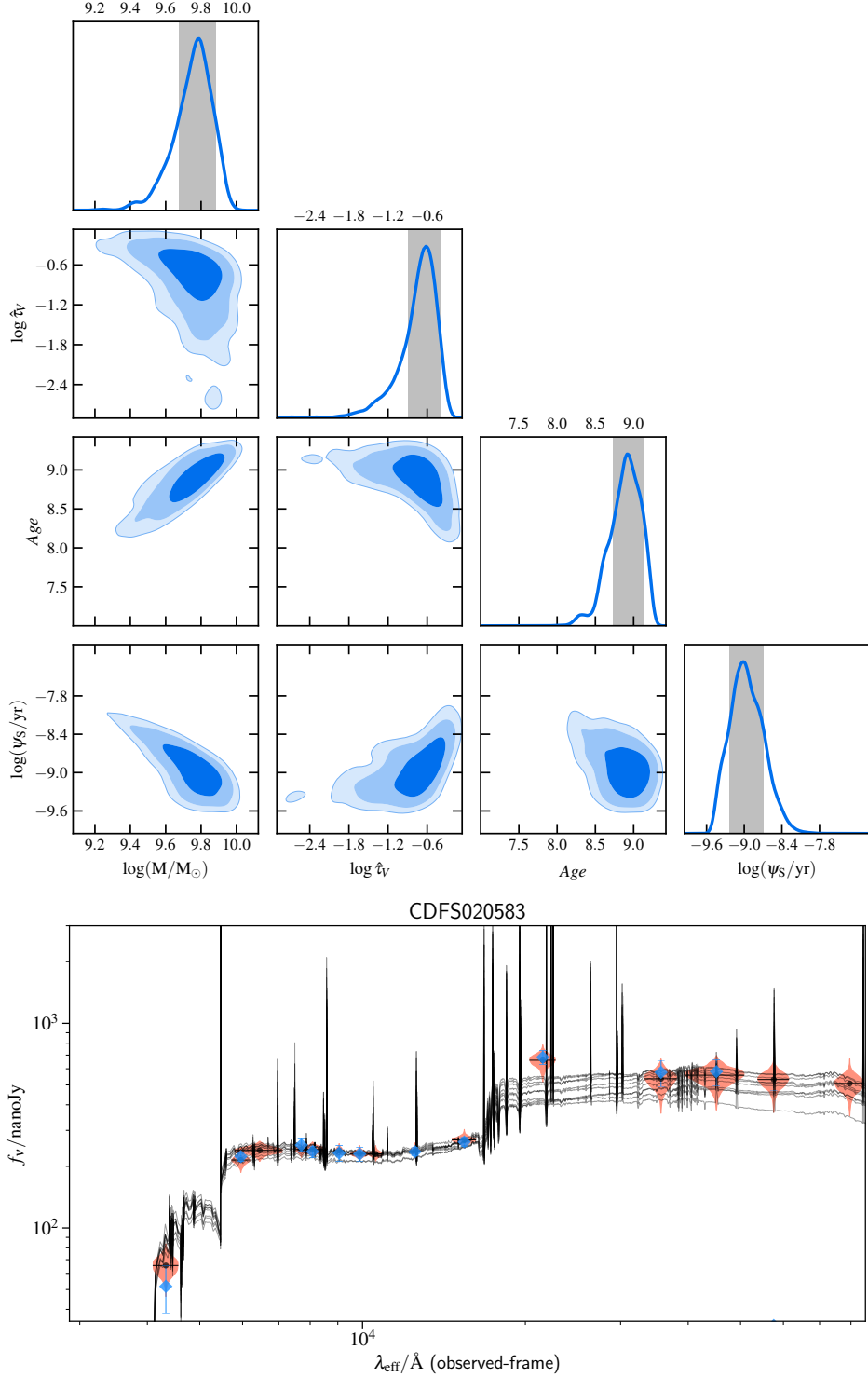


Fig. 4. *Top panel:* posterior probability distributions of the physical parameters derived for the object CDFS020583 at $z = 3.5$. From left to the right we show: the logarithmic mass in units of solar masses, the dust attenuation coefficient, the age and the logarithmic specific SFR. The different contours indicate the 1 σ , 2 σ and 3 σ confidence levels. *Bottom panel:* spectral energy distribution of the same galaxy. The blue points are the photometric observations while the orange shaded regions are the 68% confidence intervals derived by BEAGLE. The black solid lines are ten among the theoretical SEDs with the highest likelihood values.

the metallicity of the objects. Among the indicators described in Sommariva et al. (2012), we used the spectral indices F1370, F1425, F1460 and F1501, because they are not affected by the interstellar absorption lines at our resolution. The value of the standard deviation of the Gaussian prior is the dispersion of the values derived by the four indices. We remark that the aim of this

measurement is only to find an average metallicity to lower the number of free parameters in the SED fitting.

We derived the stellar masses, star formation rates (SFR), mass-weighted ages and the color excess $E(B-V)$, which measures the dust content. We show in Figure 4 an example of a SED fit of a galaxy in our sample. In the *top panel* we show the proba-

bility distribution functions of the derived physical parameters and in *bottom panel* we compare the observed photometry with the one predicted by BEAGLE. We show in Figure 5 the SFR as a function of the stellar mass for the sources in our selected sample in green and the parent sample in yellow, along with the star-formation main sequence relation evaluated by Santini et al. (2017) for star-forming galaxies in the first four *HST* Frontier Fields in the redshift range $3 \lesssim z \lesssim 4$ (blue line). We also show in the same Figure the distributions of SFRs and stellar masses for the sources in our sample (green histogram) and in the parent sample (yellow histogram). Both the samples lie on average on the star formation main sequence, as the majority of high redshift star-forming galaxies. This relation is believed to characterize galaxies that have grown on long timescales as a consequence of smooth cold gas accretion from the IGM (e.g. Dekel et al. 2009). We note that there is some scatter around the main sequence in Figure 5, that is however consistent with that of the Santini et al. (2017) sample. The reason for this scatter could be related to the SFH used in the SED fitting procedure (e.g. Cassarà et al. 2016). To check this, we have also performed the same SED-fitting using free metallicity in the range $[0.006 Z_{\odot}; 1.74 Z_{\odot}]$ and two different types of SFH: exponentially declining SFH and delayed exponentially declining SFH. In both cases the scatter around the main sequence is consistent with that obtained with constant SFH. Moreover, also the individual values are consistent. In particular, while the stellar masses are usually not affected by the choice of SFHs, the SFRs can instead vary significantly. In this case however, if we compare the results obtained with constant SFH and exponential and delayed SFHs, they are still consistent within the 1σ errors. In a future work we will use BEAGLE to assess the implications of this choice with a larger sample of galaxies in the VANDELS survey.

3.2.2. β slope

We also determined the slope β of the UV continuum. We adopted the common power-law approximation for the UV spectral range $F_{\lambda} \propto \lambda^{\beta}$ (see e.g. Castellano et al. 2012), and estimated the slope β by fitting a linear relation through the observed AB magnitudes: $M_i = -2.5(\beta + 2.0)\log\lambda_i + c$, where M_i is the magnitude in the i th filter at effective wavelength λ_i . The filters we use span the rest-frame wavelength range $\lambda \sim 1250 - 2500 \text{ \AA}$, such that the Ly α emission line is not included in any of the filters. We divided our sources in two different bins of redshifts: $3 \lesssim z \lesssim 3.5$ and $3.5 \lesssim z \lesssim 4.5$, resulting in the observed AB magnitudes I, Z, Y, and J for the first and Z, Y, J and H for the second. As for the SED fitting procedure, we used the CANDELS official photometric catalogs (Galamez et al. 2013; Guo et al. 2013) for the sources in CANDELS, and the photometric catalogs described in McLure et al. (2018) for the galaxies in the wide-field areas.

4. Parent sample versus selected sample

To determine to what extent our sample (which is mostly based on the presence of the CIII] emission line) is representative of the parent sample of Ly α emitting galaxies, we performed a Kolmogorov-Smirnov test. Since we could not check all the > 300 individual objects in the parent sample for the presence of residual of skylines or other extraction problems around the position of the Ly α emission line, we restricted the parent sample to galaxies where the fundamental parameter, the Ly α EW, was determined with an accuracy of at least 50%. We find that, in terms of stellar masses, SFRs, dust content and ages, we cannot

parameter1	parameter2	ρ	N	significance
EW(Ly α)	EW(CIII])	0.62	39	$> 3\sigma$
EW(Ly α)	age	0.40	47	$> 3\sigma$
EW(Ly α)	M*	-0.34	47	$\sim 3\sigma$
EW(Ly α)	E(B-V)	-0.43	47	$> 3\sigma$
EW(Ly α)	β	-0.56	44	$> 3\sigma$
IS _{shift}	Ly α _{shift}	0.44	23	$> 2\sigma$
EW(Ly α)	UV _{ext}	-0.37	38	$> 2\sigma$
EW(Ly α)	Ly α _{ext}	-0.39	44	$> 2\sigma$

Table 1. Spearman rank correlation coefficients for the different correlations found in this analysis sorted by the significance of the relation and the value of the Spearman rank correlation coefficient. (1) and (2) parameters for which we are testing the (anti)correlation; (3) Spearman rank correlation coefficient, ρ ; (4) number of sources in the subsets considered; (5) the significance of the relation.

reject the hypothesis that the two samples are drawn from the same distribution. For the Ly α EW distribution, we find instead a very low p-value ($p = 0.002$), meaning that our galaxies are not drawn from the same distribution of the parent sample. The difference can be clearly appreciated in Figure 6, where we show the Ly α EW distribution of the selected sample in green and that of the parent sample in yellow. The median Ly α EWs are 21.4 Å and 13.4 Å respectively. This difference was expected since we selected our sample mostly on the basis of the presence of the CIII] $\lambda 1909 \text{ \AA}$ which was necessary to evaluate the systemic redshift, and given the apparent correlation between the strength of Ly α and that of CIII] (see below). However the Ly α EW distribution spans the same range of values as the parent sample, only changing in shape. Very similar considerations can be made for the β slopes. The KS test indicates that the two samples are not drawn from the same distributions ($p = 0.06$). This is expected given that objects with larger Ly α EW have steeper β slopes (see below) and also the presence of the CIII] $\lambda 1909 \text{ \AA}$ emission line might be correlated with a harder ionizing continuum and therefore a bluer UV spectral slope. Also in this case the two samples still span the same range of β .

Because of this, along with the fact that the physical properties of the two samples are drawn from the same statistical distribution, we can reasonably conclude that the results that we find in this analysis for the selected sample will be also valid, at least in qualitative terms, for the general population of galaxies with Ly α emission. This is also corroborated by the fact that, as we show in the next Section, both samples show the same correlations between the Ly α EW and the physical properties.

5. Results

To study how the kinematics and appearance of Ly α emission depend on the galaxies physical properties and the kinematics of the ISM, we investigated the relations between the galaxy spectral and physical properties evaluated in the previous section. In particular, to give a non parametric statistical measure of the correlations between the different parameters, we evaluated the Spearman rank correlation coefficients (ρ). We show in Table 1 the Spearman rank correlation coefficients for the variables that we find (anti)correlated, and the significance of these relationships. We trust the correlation only if the significance is higher than 2σ .

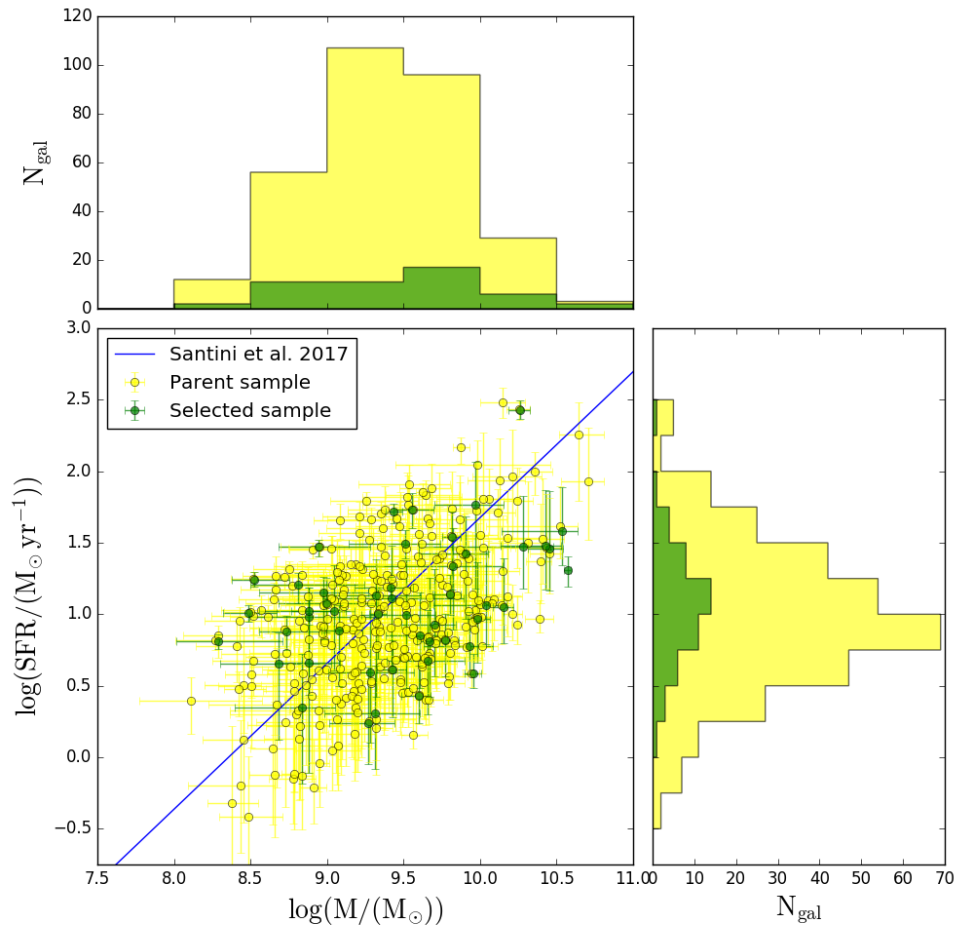


Fig. 5. SFR as a function of the stellar mass for the objects in our sample (green points) and the parent sample (yellow points). The blue line is one of the latest star formation main sequence evaluated by Santini et al. (2017) for galaxies in the first four *HST* Frontier Fields in the redshift range $3 \lesssim z \lesssim 4$.

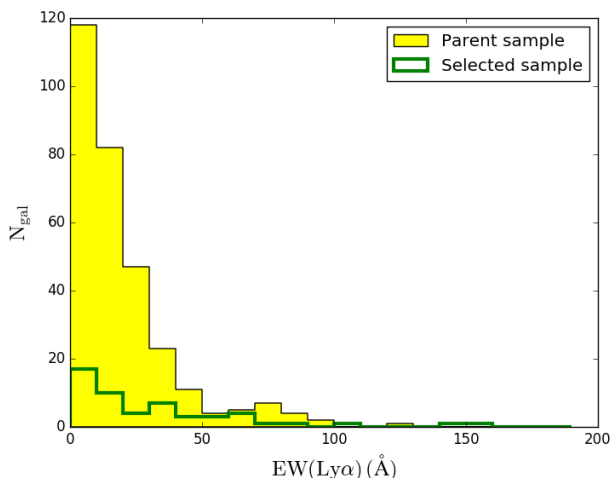


Fig. 6. Ly α EW distributions for the sources in the *parent sample* (yellow histogram) and those in the *selected sample* (green empty histogram).

5.1. Correlations between Ly α EW and physical properties

In this subsection we discuss the main correlations that we find between the Ly α EW and the galaxies' physical properties. We first discuss the results obtained for the selected sample, then we compare them with the corresponding relations for the parent sample.

– EW(Ly α) vs M*

We show in Figure 7 (blue points) the relation between EW(Ly α) and stellar mass M*: the two parameters are moderately anti-correlated with a Spearman rank correlation coefficient of -0.34 . While the lack of objects with large masses and large Ly α EW is surely real, we might be losing objects with small masses and small EWs (lower left part of the figure) because of the flux limit of the survey. We checked that this not the case, given the sensitivity of our observations to emission lines and the median R band magnitude of our targets.

The existence of this relation is quite controversial since it has been previously observed by some works (e.g. Finkelstein et al. 2007; Jones et al. 2012; Hathi et al. 2016) but not by others (e.g. Kornei et al. 2010). In particular Pentericci et al. (2007, 2010) found that galaxies without Ly α in emission tend to be more massive and dustier than the rest of the LBGs sample and found a significant lack of massive galaxies with high EW(Ly α), which could be explained if the most

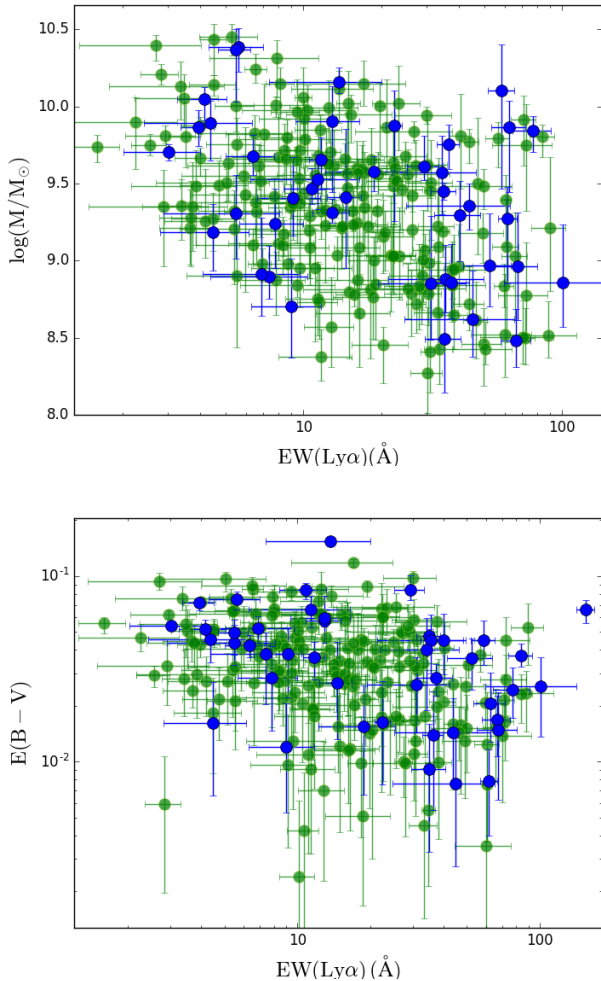


Fig. 7. Correlations between the Ly α EW and the stellar mass (top panel) and between the Ly α EW and the color excess E(B-V) (bottom panel) for the selected sample (blue) and the parent sample (green).

massive galaxies were either dustier and/or if they contained more neutral gas than less massive objects. However they did not find a precise one-to-one correlation. Finally, Du et al. (2018) found a clear relation for stacks of spectroscopically confirmed LBGs at $z \sim 4$ but they showed that this becomes weaker as redshift decreases with a almost flat trend at $z \sim 3$.

– EW(Ly α) vs E(B-V)

As shown in the bottom panel of Figure 7, we also observe a moderate anti-correlation between dust extinction, indicated by the color excess E(B-V), and EW(Ly α) with a Spearman rank correlation coefficient of -0.43 , in agreement with several previous studies (Shapley et al. 2003; Kornei et al. 2010; Pentericci et al. 2010; Jones et al. 2012; Erb et al. 2016; Hathi et al. 2016). This is also consistent with theoretical expectations: Ly α emission is easily quenched by dust, since Ly α photons can be absorbed by dust grains, that then re-emit them as thermal emission.

– EW(Ly α) vs galaxy age

We observe a correlation between Ly α EW and the mass-weighted galaxy age, in the sense that galaxies with a higher Ly α EW, are older than galaxies with lower Ly α EW. The Spearman rank correlation coefficient is 0.40 . While the relation between Ly α and E(B-V) is now quite well established, the relation between Ly α and galaxy age is still highly de-

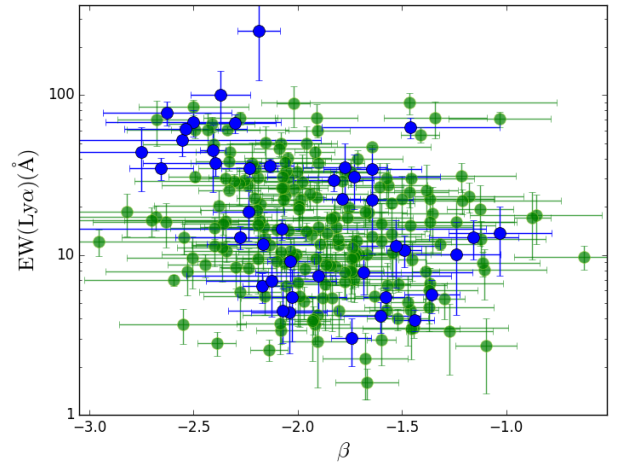


Fig. 8. Correlation between the β slope and the Ly α EW.

bated. Some studies claim that Ly α emitting galaxies are mostly young galaxies in the very early stages of formation (e.g. Pentericci et al. 2007; Du et al. 2018), while others claim that Ly α emission is typically found in older galaxies (e.g. Shapley et al. 2003; Kornei et al. 2010). Finally, other studies do not find apparent relation between Ly α EW and age (e.g. Pentericci et al. 2009, 2010).

According to our observations, the brightest Ly α emitters are also the oldest. This could be explained by a scenario in which the oldest galaxies have experienced in the past strong burst of star formation that caused massive stellar winds with the consequent expulsion of neutral gas and dust, resulting in a reduced covering fraction and therefore a smaller attenuation of the Ly α photons, while the younger galaxies still have more dust and therefore smaller Ly α EWs.

We remark that the ages are one of the most uncertain parameters in SED fitting also because they are somehow dependent on the SFH adopted (see e.g. Carnall et al. 2018; Leja et al. 2018): we tested the exponentially declining SFH and the delayed exponentially declining SFH instead of a constant one, but we do not find significant differences in the observed correlation.

– EW(Ly α) vs SFR

We do not find any correlation between the Ly α EW and the SFR, measuring a Spearman rank correlation coefficient of 0.14 , with very high confidence. This might look in contrast with the results from several previous studies (e.g. Kornei et al. 2010; Hathi et al. 2016) that found instead that these two quantities are anti-correlated, with the brightest LAEs having lower SFRs. However, all these studies analysed galaxies with Ly α both in absorption and in emission, while we have only selected galaxies with Ly α emission. Indeed, if we restrict previous results to only those sources with Ly α in emission, no clear correlations are observed. This suggests that galaxies with Ly α emission have on average smaller SFRs if compared to the non emitting galaxies, but that there is no particular trend with the actual strength of Ly α . Note also that while the sources presented in Hathi et al. (2016) are in the same mass range as those analysed in this paper, the sources in Kornei et al. (2010) have much larger average stellar masses ($\log(M/M_\odot) \sim 9.9$).

– EW(Ly α) vs β slope

We reproduce the well known anti-correlation between the UV spectral slope β and the Ly α EW, with a Spearman rank

correlation coefficient of -0.44 and high significance ($>3\sigma$). This relation is shown in Figure 8 for the sources in the selected sample in blue and for those in the parent sample in green.

The UV β slope depends on several parameters including dust, age and metallicity, although it is thought to be primarily driven by the dust content (e.g. Reddy et al. 2010; Hathi et al. 2016). The anti-correlation between Ly α EW and the β slope, is therefore a direct consequence of the fact that less dusty galaxies tend to show brighter Ly α emission as also probed by the anti-correlations that we observe between the Ly α EW with $E(B-V)$.

We finally tested if the relations between $EW(Ly\alpha)$ and the physical properties are also valid for the parent sample, that contains galaxies with similar physical properties but with different distributions of Ly α EWs and β slopes. The results of this analysis are also shown in Figure 7. The correlation between $EW(Ly\alpha)$ and stellar mass and $E(B-V)$ are comparable to those found for the selected sample, and similarly there is no relation between SFR and $EW(Ly\alpha)$. We however do not observe any correlation between the Ly α EW and the galaxy mass-weighted age for the galaxies in the parent sample, that was instead present for the sources in the selected sample. We remark however that, as already mentioned above, the galaxy age is the most uncertain parameter from SED fitting and that the Spearman rank correlation coefficients and their significance, are evaluated without taking into account the uncertainty in the individual parameters. Finally, a correlation between the Ly α EW and β slope is observed also for the parent sample, even if it is weaker than that observed for the selected sample ($\rho = 0.2$ with $> 2\sigma$ significance).

5.2. Correlations between Ly α EW and spectral properties

We discuss in this subsection the relations obtained between the Ly α EW and the other galaxy spectral properties.

– $EW(Ly\alpha)$ vs $EW(CIII)]$

We observe a strong correlation between $EW(Ly\alpha)$ and $EW(CIII)]$, with Spearman rank correlation coefficient of 0.62 , that we show in the top panel of Figure 9. This relation has been previously investigated by other studies, using both stacks and individual spectra with somewhat contradictory results. Shapley et al. (2003) found a positive trend between the two quantities on spectral stacks of a sample of LBGs at $z \sim 3$. More recently Stark et al. (2014) found a tighter relation between $EW(Ly\alpha)$ and $EW(CIII)]$ in individual galaxies in a sample of 17 strongly lensed low-mass galaxies at $z \sim 2$. In contrast to these results, other studies just show tentative correlations with a very large scatter. For instance, Rigby et al. (2015) studied a sample of galaxies at $1.6 < z < 7$, and found an apparent correlation, though with appreciable scatter especially for the weakest Ly α and CIII] emitters ($EW(Ly\alpha) \lesssim 50\text{\AA}$ and $EW(CIII)] \lesssim 5\text{\AA}$). A very similar large scatter was found by Le Fèvre et al. (2017) for individual star-forming galaxies at $2 < z < 3.8$ with CIII] emission in VUDS, with a clear trend that is only observed using stacks as in Shapley et al. (2003). Finally, we mention that Du et al. (2018), using stacks of $z \sim 2$ star-forming galaxies observed a nearly null correlation between $EW(CIII)]$ and $EW(Ly\alpha)$, with the exception of the strongest Ly α bin where they observe a higher $EW(CIII)]$ with respect to the other sub-samples. All these works suggest that, on average, the presence of Ly α (CIII]) emission increases the

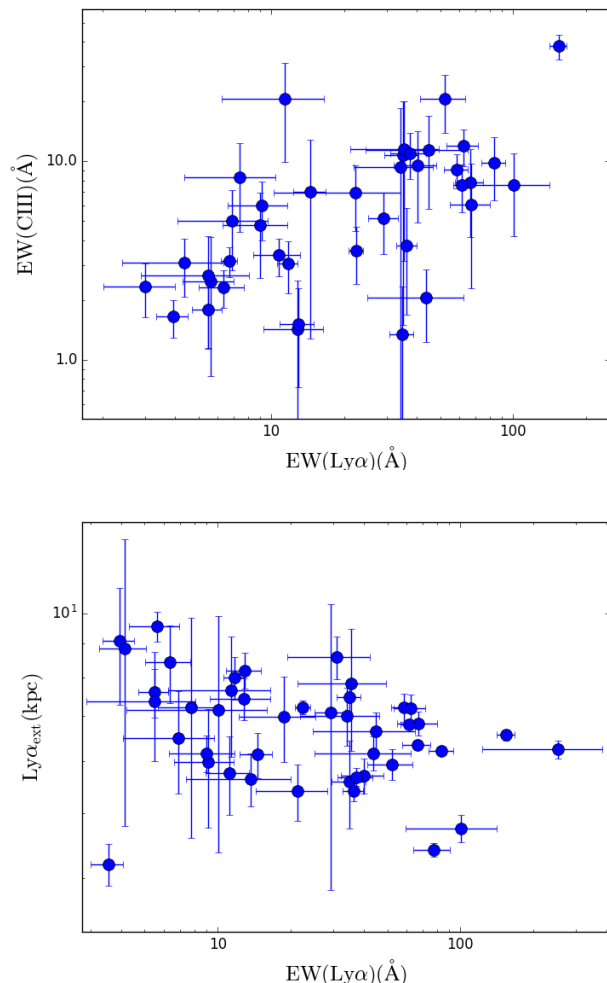


Fig. 9. Correlations between the Ly α EW and the CIII] EW (top panel) and between the Ly α EW and the Ly α spatial extent (bottom panel).

probability of detecting also CIII] (Ly α) emission, but that there is not a clear one to one relation. There are indeed numerous examples of galaxies with CIII] emission that show only a weak Ly α emission or no emission at all, and several LAEs that do not show any sign of CIII] $\lambda 1909\text{\AA}$ emission line (e.g. Le Fèvre et al. 2017). This behavior is not surprising. The production of both Ly α and CIII] photons strongly depends indeed on the on-going star formation but there are significant differences in the production and radiative transfer processes of the two lines (see e.g. Nakajima et al. 2018, for the CIII] production). In our data we find a rather strong correlation between the Ly α EW and CIII] EW in individual galaxies.⁵ We note that our sample contains galaxies with higher stellar masses compared to e.g. Stark et al. (2014) who probe a mass range of $9 \times 10^6 - 1.3 \times 10^9 M_\odot$. Clearly the tight relation that we observe is somehow biased since we pre-selected the sample to have both Ly α and CIII] emission. We are therefore missing the objects with CIII] (Ly α) emission and not Ly α (CIII]). Indeed a similar result was also found by Amorín et al. (2017) for a sample of 10 low-mass galaxies at $z \sim 3$, selected by their strong Ly α emission ($EW(Ly\alpha) > 45\text{\AA}$) in VUDS. In a future work we will ex-

⁵ The correlation is unchanged if we exclude the point in the upper right part of Figure 9, that might appear to drive it.

plot the whole VANDELS dataset to evaluate the unbiased relation between $\text{Ly}\alpha$ and CIII].

– EW($\text{Ly}\alpha$) vs $\text{Ly}\alpha_{\text{ext}}$ and UV $_{\text{ext}}$

We observe a mild anti-correlation between EW($\text{Ly}\alpha$) and the UV and $\text{Ly}\alpha$ spatial extents, in the sense that galaxies with higher EW appear to be more compact in size both in the UV continuum and in $\text{Ly}\alpha$. The Spearman rank test gives correlation coefficients of -0.37 and -0.39, respectively. The relation between $\text{Ly}\alpha$ EW and UV size has been previously observed in Lyman Break galaxies (e.g. Jones et al. 2012). The relation with $\text{Ly}\alpha_{\text{ext}}$ has been previously observed by Guaita et al. (2017), using stacks of VUDS galaxies at $z \sim 3$, although their definition of $\text{Ly}\alpha$ extension was slightly different than ours since they subtracted in quadrature to the $\text{Ly}\alpha_{\text{ext}}$, the UV continuum to implicitly consider the deconvolution with the PSF of the observations. We now find a similar correlation also using individual objects.

We tested these correlations also on a small sub-sample of 17 galaxies, for which we could measure all the quantities listed in Section 3. We find for this sample the same correlations observed before and with the same significance.

5.3. Correlations with IS shift

We tested if the velocity of the IS medium is related to other galaxy parameters. The only property that we find related to the IS_{shift} , is the offset of the $\text{Ly}\alpha$ line with respect to the systemic redshift, that we show in Figure 10, along with the distributions of the two parameters, in yellow and green, respectively.

The Spearman rank correlation coefficient for this relation is 0.44, with a significance that is $> 2\sigma$. This relation has been previously observed by Guaita et al. (2017) using stacks of $z \sim 3$ star forming galaxies. We discuss this trend and its interpretation in the next Section.

We do not observe any correlation between IS shift and other physical parameters. In particular while we expected some trend with the SFR, the Spearman rank correlation coefficient is only -0.15, with a significance that is $< 2\sigma$. This trend is well established in low redshift samples (e.g. Weiner et al. 2009; Heckman et al. 2015) but not at high redshift. In principle, a relation would be expected also at high redshift since outflows are believed to be driven by pressure created by supernovae, stellar winds, and the radiation field caused by star formation. However our SFRs are derived from SED fitting and have uncertainties: to properly test this relation, SFR derived from $\text{H}\alpha$ line flux would be needed. Besides SFR, we also specifically tested the correlation between outflow velocity and SFR surface density, as defined in Heckman et al. (2011), and between outflow velocity and specific star formation rate defined as SFR/M^* but again we find no significant trend in our sample.

6. Discussion: the origin of the relation between ISM and $\text{Ly}\alpha$ velocity shifts

The correlation observed in Figure 10 is the most interesting result that we find in this analysis since it relates the $\text{Ly}\alpha$ line kinematics to that of the ISM. According to theoretical predictions, the $\text{Ly}\alpha$ kinematics depends on the velocity of the ISM, but also on the HI column density, that gives a measure of the amount of neutral gas along the line of sight. In this Section we interpret the relation found between $\text{Ly}\alpha$ velocity and ISM outflow in the context of the theoretical model of Verhamme et al.

(2006), that accurately predicts the $\text{Ly}\alpha$ shift, $\text{Ly}\alpha$ EW and spatial extension as functions of the ISM kinematics and HI column density. In this framework galaxies are modeled as a central region of young stars surrounded by an expanding, spherical, homogeneous, and iso-thermal shell of neutral hydrogen. In this scenario, small $\text{Ly}\alpha$ shifts and large (but less than $\sim 300 \text{ km/s}$) outflow velocities are predicted only for small HI column densities ($\sim 10^{19} \text{ cm}^{-2}$), while large $\text{Ly}\alpha$ shifts but with small or null outflow velocities can be observed only in case of large HI column densities ($\sim 10^{20} - 10^{21} \text{ cm}^{-2}$) because the $\text{Ly}\alpha$ photons must have undergone many scatters before escaping the galaxy, resulting in higher $\text{Ly}\alpha$ velocity offsets with respect to the systemic redshift. According to this, we expect the objects in the lower left part of Figure 10 to be characterized by smaller HI column densities, while the objects in the upper right part of the plot, by larger HI column densities.

This scenario is also consistent with the distribution of the $\text{Ly}\alpha$ spatial extension with respect to the UV continuum of our sources. We show in Figure 11 the same plot as in Figure 10, but color coded for the $\text{Ly}\alpha$ spatial extent with respect to the UV extension ($\text{Ly}\alpha_{\text{ext}}/\text{UV}_{\text{ext}}$). The objects that appear to have lower HI column densities (from the $\text{Ly}\alpha$ vs IS kinematics) also show on average more compact $\text{Ly}\alpha$ spatial profile, while the objects that should have larger HI column density (from the $\text{Ly}\alpha$ kinematics) also have large $\text{Ly}\alpha$ extension compared to the UV extensions. These observations are perfectly in agreement with the shell model predictions, since broader $\text{Ly}\alpha$ spatial profile with respect to the continuum are expected if the $\text{Ly}\alpha$ photons undergo many scatters before escaping the galaxy (caused by larger HI column densities), while more compact spatial profile are predicted for galaxies with lower HI column densities, since the $\text{Ly}\alpha$ photons can escape galaxies in a more direct way (Verhamme et al. in preparation). Assuming that this model is actually working, and that the trend observed in Figure 10, is due to a different HI column density, we then try to assess if there is any apparent relation between $\text{N}(\text{HI})$ and other galaxy physical properties. In particular, we show in Figure 12, the same relation between IS_{shift} and $\text{Ly}\alpha_{\text{shift}}$ color coded for the stellar mass, the attenuation of the dust, indicated by the color excess $E(B-V)$, the SFR and the $\text{Ly}\alpha$ EW. None of these physical parameters appears to correlate with the HI column density, i.e. we do not find any significant change of physical properties in the galaxies as we move from the lower-left part of the figure (small HI column densities) to the upper-right part (high HI column densities). A possible limitation of this analysis is that we are probing only a relatively small range for these quantities. For example our galaxies span only a factor of 10 in mass: it is possible that if a weak correlation between $\text{N}(\text{HI})$ and mass exist, it would emerge only by probing much larger ranges of values.

We also do not observe any trend between $\text{N}(\text{HI})$ and $\text{Ly}\alpha$ EW. A tentative anti-correlation between these quantities has been observed in local galaxies (e.g. Pardy et al. 2014; Yang et al. 2017). From a theoretical point of view, this would be expected only if the intrinsic EW and/or the gas to dust ratio were approximately constant. In the first case, with the increase of $\text{N}(\text{HI})$, the path length of $\text{Ly}\alpha$ photons would increase and so their chance of being absorbed by dust, while in the second case an increase in $\text{N}(\text{HI})$ would also lead to an increase in the dust optical depth, suppressing the $\text{Ly}\alpha$ EW. We do not observe this relation. This is however not in contrast with our results since we do not have reasons to believe that the conditions to observe this relation are met in our sample of galaxies. From Figure 12 we do not observe indeed any relation between dust attenuation and the HI column density, suggesting a non constant gas to dust ratio. Moreover,

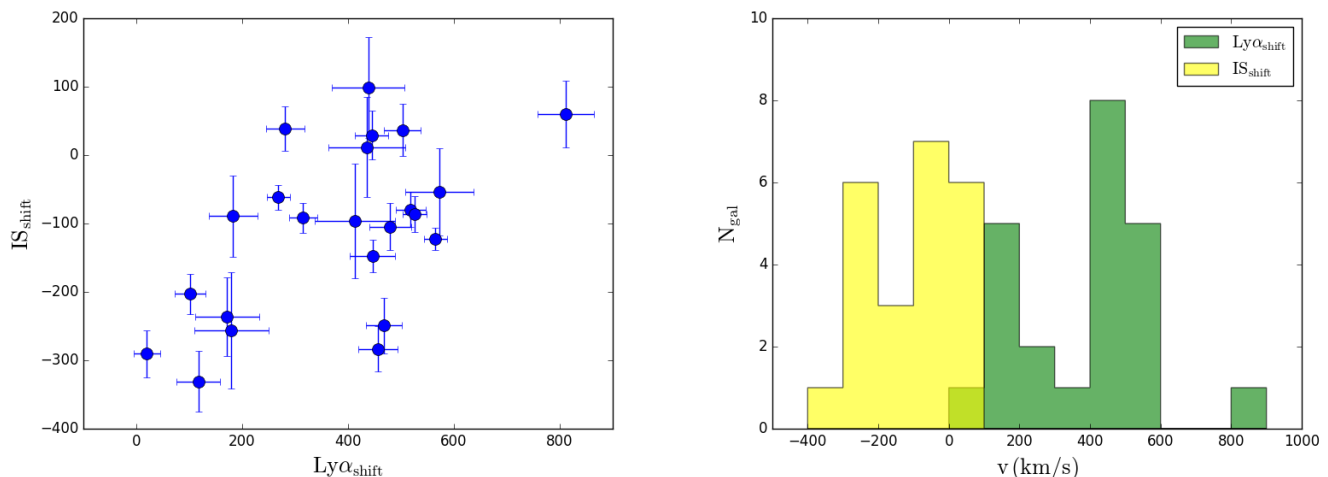


Fig. 10. *Left panel:* anti-correlation between the velocity of the main peak of the Ly α profile with respect to the systemic redshift and the IS velocity shift. *Right panel:* Histogram of the measured (centroid) velocities of interstellar absorption lines in yellow and Ly α emission in green with respect to the galaxy systemic redshift derived from the centroid of either the CIII] λ 1909Å or the HeII λ 1640Å emission lines. The yellow sample contains only the 27 objects for which we could give a precise measure of the IS velocity shifts, and the green sample contains the 46 spectra where we could estimate the Ly α peak shift. The median values of the velocity offsets are $IS_{shift} = -89$ km/s and $Ly\alpha_{shift} = 400$ km/s, respectively.

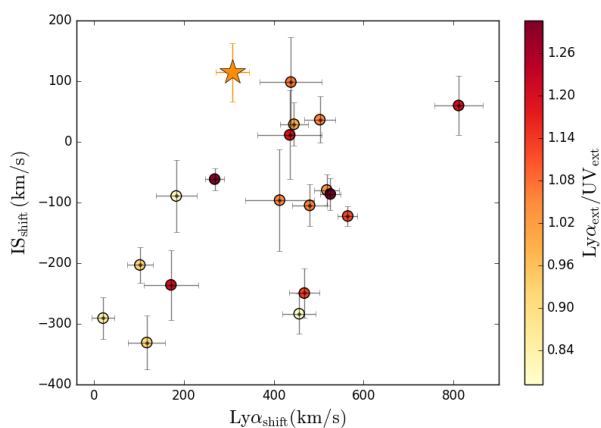


Fig. 11. Ly α velocity offset vs IS velocity offset with respect to the systemic redshift. The plot is color coded for the Ly α spatial extension with respect to the UV continuum, as shown from the color bar in the right side of the Figure. The star on the plot is the stack of the sources with no individual IS_{shift} measure. Its colour is given by the median of the measured $Ly\alpha_{ext}/UV_{ext}$ of the IS undetected sample.

we do not observe any correlation between the Ly α peak shift and $EW(Ly\alpha)$, relation that would be expected only if the Ly α EW was correlated with $N(HI)$, since larger HI column densities cause more scattering of the Ly α photons, shifting the line peak to redder wavelengths.

What we actually observe from Figure 12 is that galaxies with brighter Ly α emission show on average no outflows and Ly α peak shifts around 400 km/s. It is not yet clear if there is a real physical motivation under these observations and this is still under investigation.

7. Summary and conclusions

In this work we have investigated the correlations between the physical properties, the Ly α properties and the kinematics of

the inter-stellar medium in a sample of 52 star-forming galaxies selected to have Ly α emission and systemic redshift from the VANDELS survey. We find that the Ly α EW and the CIII] λ 1909Å EW are strongly related, in agreement with previous observational studies at similar redshifts (e.g. Stark et al. 2014), suggesting an intrinsic connection between the production of the two lines. We note however that for this correlation the sample is biased, since we only selected galaxies with both Ly α and CIII] λ 1909Å. This relation could therefore change if the entire population of galaxies with CIII] λ 1909Å emission would be considered. We also observe a good anti-correlation, although with some scatter, between the Ly α EW and the UV and Ly α spatial extents, with compact sources being the brightest Ly α emitters. The Ly α line appears also brighter in lower mass and less dusty galaxies, while no correlation appear to exist with the SFR. We finally observe a very interesting correlation between the Ly α velocity shift and IS shift in the sense that galaxies with large ISM outflow velocities also show small Ly α velocity shifts.

According to the shell model of Verhamme et al. (2006), this relation can be explained if the galaxies with small Ly α shifts also have low HI column densities: indeed looking at their morphology, these galaxies also show compact Ly α_{ext}/UV_{ext} , in agreement with the expectations. Compact Ly α spatial profile with respect to the continuum are indeed expected for galaxies with small $N(HI)$, because the Ly α photons did not undergo many scatters with the neutral gas while escaping the galaxies (Verhamme et al. in preparation). On the other hand, objects with no ISM outflows show large Ly α shifts (~ 500 km/s). In the context of the shell model this should be observed only for galaxies which have large HI column densities ($\sim 10^{20} - 10^{21} \text{ cm}^{-2}$). Indeed these galaxies also show large $Ly\alpha_{ext}/UV_{ext}$, in agreement with the expectations since in these cases the Ly α photons scatter many times before escaping, resulting in broader spatial profiles.

We finally do not observe any correlation between the trend observed for the IS and Ly α shifts (and therefore $N(HI)$) and any other physical properties such as $EW(Ly\alpha)$, SFR, Mass, $sSFR$, Σ_{SFR} and dust content.

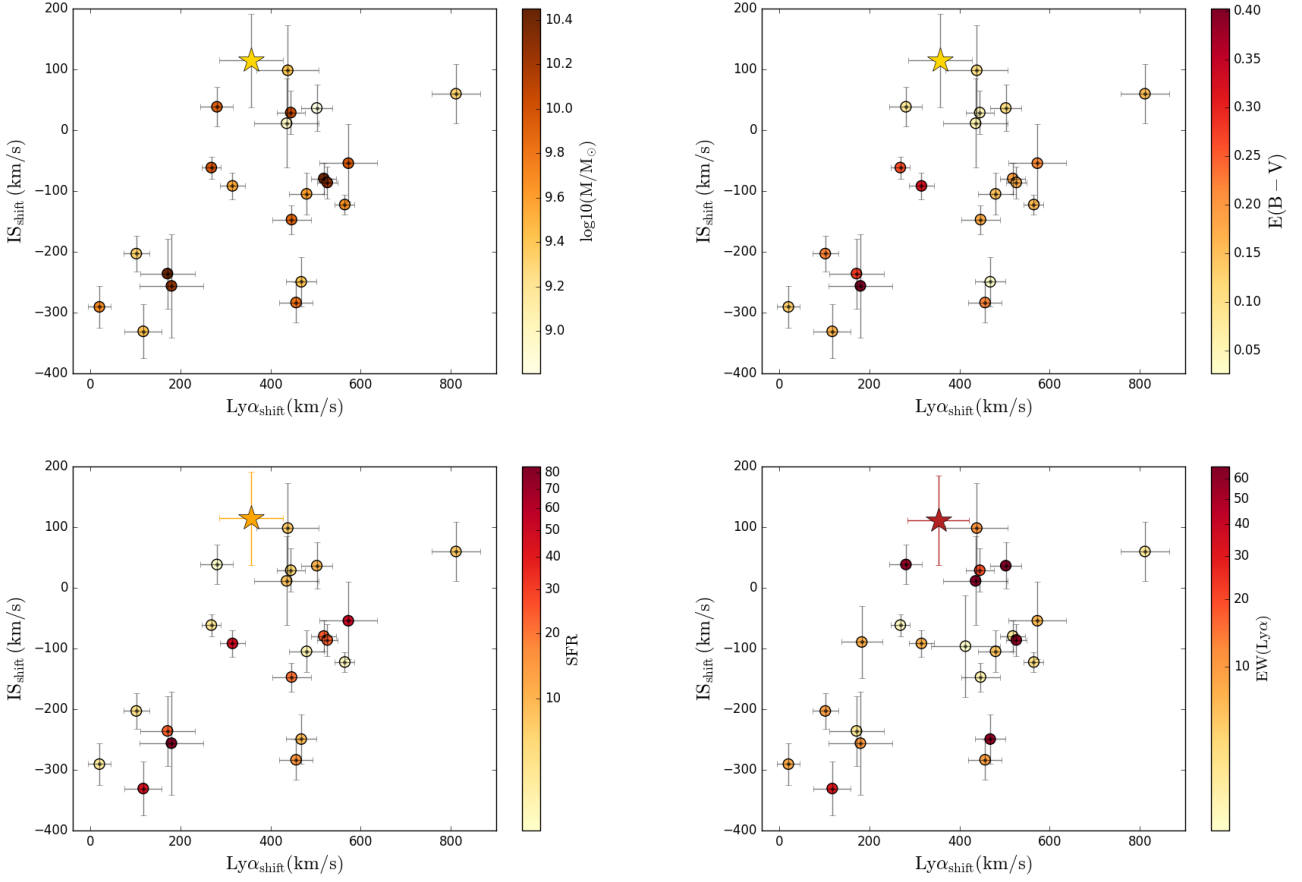


Fig. 12. IS_{shift} vs $Ly\alpha_{\text{shift}}$ color coded for different galaxy properties. From left to the right and top to bottom: stellar mass, $E(B-V)$ color excess, SFR and $EW(Ly\alpha)$. The star on each plot is the stack of the sources with no individual IS_{shift} measure, with the colour given by the median values of the measured quantity that is indicated in the color bar, for the galaxies in the IS undetected sample.

To better understand these correlations and assess their validity, we are planning to enlarge the analysis carried out so far to the entire VANDELS data-set, which will be available in the near future, since the present work was applied only to half of the final VANDELS sample.

We finally remark that the sources presented in this work have been selected as LBGs and not as LAEs, being therefore on average brighter in the UV continuum than the objects selected with a narrow band technique that are tuned to the $Ly\alpha$. This means that we are not probing the least massive sources, that are thought to be the most common galaxies in the very early Universe and possibly also those that contributed mostly to the reionization photon budget. The correlations that we find could therefore be stronger (or different) if investigated in a samples which could also include these lower mass galaxies. We are planning to test this hypothesis, by collecting a sample of strong LAEs, selected based on narrow band studies, to investigate how they compare with the VANDELS galaxies in the present sample and if there are significant differences in the correlations that we observe based on the samples selection.

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References

- Amorín, R., Fontana, A., Pérez-Montero, E., et al. 2017, *Nature Astronomy*, 1, 0052
- Bottini, D., Garilli, B., Maccagni, D., et al. 2005, *PASP*, 117, 996
- Calzetti, D., Armus, L., Bohlin, R. C., et al. 2000, *ApJ*, 533, 682
- Carnall, A. C., Leja, J., Johnson, B. D., et al. 2018, *arXiv e-prints* [arXiv:1811.03635]
- Cassarà, L. P., Maccagni, D., Garilli, B., et al. 2016, *A&A*, 593, A9
- Cassata, P., Le Fèvre, O., Charlot, S., et al. 2013, *A&A*, 556, A68
- Castellano, M., Fontana, A., Grazian, A., et al. 2012, *A&A*, 540, A39
- Chabrier, G. 2003, *PASP*, 115, 763
- Chevallard, J. & Charlot, S. 2016, *MNRAS*, 462, 1415
- Cowie, L. L., Barger, A. J., & Hu, E. M. 2011, *ApJ*, 738, 136
- Cullen, F., McLure, R. J., Khochfar, S., et al. 2018, *MNRAS*, 476, 3218
- Dekel, A., Birnboim, Y., Engel, G., et al. 2009, *Nature*, 457, 451
- Du, X., Shapley, A. E., Reddy, N. A., et al. 2018, *ApJ*, 860, 75
- Ebbets, D. 1995, in *Calibrating Hubble Space Telescope. Post Servicing Mission*, ed. A. P. Koratkar & C. Leitherer, 207
- Erb, D. K., Pettini, M., Steidel, C. C., et al. 2016, *ApJ*, 830, 52
- Finkelstein, K. D., Finkelstein, S. L., Tilvi, V., et al. 2015, *ApJ*, 813, 78
- Finkelstein, S. L., Rhoads, J. E., Malhotra, S., & Grogin, N. 2009, *ApJ*, 691, 465
- Finkelstein, S. L., Rhoads, J. E., Malhotra, S., Grogin, N., & Wang, J. 2008, *ApJ*, 678, 655
- Finkelstein, S. L., Rhoads, J. E., Malhotra, S., Pirzkal, N., & Wang, J. 2007, *ApJ*, 660, 1023
- Galametz, A., Grazian, A., Fontana, A., et al. 2013, *ApJS*, 206, 10
- Garilli, B., Fumana, M., Franzetti, P., et al. 2010, *PASP*, 122, 827
- Garilli, B., Paiero, L., Scodreggio, M., et al. 2012, *PASP*, 124, 1232
- Grogin, N. A., Kocevski, D. D., Faber, S. M., et al. 2011, *ApJS*, 197, 35
- Gronke, M. & Dijkstra, M. 2016, *ApJ*, 826, 14
- Guaita, L., Talia, M., Pentericci, L., et al. 2017, *A&A*, 606, A19
- Guo, Y., Ferguson, H. C., Giavalisco, M., et al. 2013, *ApJS*, 207, 24
- Gutkin, J., Charlot, S., & Bruzual, G. 2016, *MNRAS*, 462, 1757
- Hathi, N. P., Le Fèvre, O., Ilbert, O., et al. 2016, *A&A*, 588, A26

- Heckman, T. M., Alexandroff, R. M., Borthakur, S., Overzier, R., & Leitherer, C. 2015, *ApJ*, 809, 147
- Heckman, T. M., Borthakur, S., Overzier, R., et al. 2011, *ApJ*, 730, 5
- Jones, T., Stark, D. P., & Ellis, R. S. 2012, *ApJ*, 751, 51
- Koekemoer, A. M., Faber, S. M., Ferguson, H. C., et al. 2011, *ApJS*, 197, 36
- Kornei, K. A., Shapley, A. E., Erb, D. K., et al. 2010, *ApJ*, 711, 693
- Le Fèvre, O., Lemaux, B. C., Nakajima, K., et al. 2017, *ArXiv e-prints* [arXiv:1710.10715]
- Le Fèvre, O., Tasca, L. A. M., Cassata, P., et al. 2015, *A&A*, 576, A79
- Leja, J., Carnall, A. C., Johnson, B. D., Conroy, C., & Speagle, J. S. 2018, *arXiv e-prints* [arXiv:1811.03637]
- Lenz, D. D. & Ayres, T. R. 1992, *PASP*, 104, 1104
- Marchi, F., Pentericci, L., Guaita, L., et al. 2018, *A&A*, 614, A11
- Maseda, M. V., Brinchmann, J., Franx, M., et al. 2017, *A&A*, 608, A4
- Matthee, J., Sobral, D., Gronke, M., et al. 2018, *A&A*, 619, A136
- McLure, R. J., Pentericci, L., Cimatti, A., et al. 2018, *MNRAS*, 479, 25
- Nakajima, K., Schaerer, D., Le Fèvre, O., et al. 2018, *A&A*, 612, A94
- Osterbrock, D. E. & Ferland, G. J. 2006, *Astrophysics of gaseous nebulae and active galactic nuclei*
- Pardy, S. A., Cannon, J. M., Östlin, G., et al. 2014, *ApJ*, 794, 101
- Pentericci, L., Grazian, A., Fontana, A., et al. 2009, *A&A*, 494, 553
- Pentericci, L., Grazian, A., Fontana, A., et al. 2007, *A&A*, 471, 433
- Pentericci, L., Grazian, A., Scarlata, C., et al. 2010, *A&A*, 514, A64
- Pentericci, L., McLure, R. J., Garilli, B., et al. 2018, *A&A*, 616, A174
- Quider, A. M., Pettini, M., Shapley, A. E., & Steidel, C. C. 2009, *MNRAS*, 398, 1263
- Reddy, N. A., Erb, D. K., Pettini, M., Steidel, C. C., & Shapley, A. E. 2010, *ApJ*, 712, 1070
- Reddy, N. A., Steidel, C. C., Fadda, D., et al. 2006, *ApJ*, 644, 792
- Rigby, J. R., Bayliss, M. B., Gladders, M. D., et al. 2015, *ApJ*, 814, L6
- Santini, P., Fontana, A., Castellano, M., et al. 2017, *ApJ*, 847, 76
- Schaerer, D., Fragos, T., & Izotov, Y. I. 2019, *A&A*, 622, L10
- Scoddeggio, M., Franzetti, P., Garilli, B., et al. 2005, *PASP*, 117, 1284
- Shapley, A. E., Steidel, C. C., Pettini, M., & Adelberger, K. L. 2003, *ApJ*, 588, 65
- Shimasaku, K., Kashikawa, N., Doi, M., et al. 2006, *PASJ*, 58, 313
- Sommariva, V., Mannucci, F., Cresci, G., et al. 2012, *A&A*, 539, A136
- Stark, D. P., Ellis, R. S., Charlot, S., et al. 2017, *MNRAS*, 464, 469
- Stark, D. P., Ellis, R. S., Chiu, K., Ouchi, M., & Bunker, A. 2010, *MNRAS*, 408, 1628
- Stark, D. P., Richard, J., Siana, B., et al. 2014, *MNRAS*, 445, 3200
- Steidel, C. C., Bogosavlevic, M., Shapley, A. E., et al. 2018, *ArXiv e-prints* [arXiv:1805.06071]
- Talia, M., Mignoli, M., Cimatti, A., et al. 2012, *A&A*, 539, A61
- Verhamme, A., Orlitová, I., Schaerer, D., & Hayes, M. 2015, *A&A*, 578, A7
- Verhamme, A., Orlitová, I., Schaerer, D., et al. 2017, *A&A*, 597, A13
- Verhamme, A., Schaerer, D., & Maselli, A. 2006, *A&A*, 460, 397
- Weiner, B. J., Coil, A. L., Prochaska, J. X., et al. 2009, *ApJ*, 692, 187
- Yamada, T., Matsuda, Y., Kousai, K., et al. 2012, *ApJ*, 751, 29
- Yang, H., Malhotra, S., Gronke, M., et al. 2017, *ArXiv e-prints* [arXiv:1701.01857]